

# GAS DISTRIBUTION, KINEMATICS, AND EXCITATION STRUCTURE IN THE DISKS AROUND THE CLASSICAL BE STARS $\beta$ CANIS MINORIS AND $\zeta$ TAURI<sup>1</sup>

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## ABSTRACT

Using CHARA and VLTI near-infrared spectro-interferometry with hectometric baseline lengths (up to 330 m) and with high spectral resolution (up to  $\lambda/\Delta\lambda = 12\,000$ ), we studied the gas distribution and kinematics around two classical Be stars. The combination of high spatial and spectral resolution achieved allows us to constrain the gas velocity field on scales of a few stellar radii and to obtain, for the first time in optical interferometry, a dynamical mass estimate using the position-velocity analysis technique known from radio astronomy. For our first target star,  $\beta$  Canis Minoris, we model the H+K-band continuum and Br $\gamma$ -line geometry with a near-critical rotating stellar photosphere and a geometrically thin equatorial disk. Testing different disk rotation laws, we find that the disk is in Keplerian rotation ( $v(r) \propto r^{-0.5 \pm 0.1}$ ) and derive the disk position angle ( $140 \pm 1.7^\circ$ ) inclination ( $38.5 \pm 1^\circ$ ), and the mass of the central star ( $3.5 \pm 0.2 M_\odot$ ). As a second target star, we observed the prototypical Be star  $\zeta$  Tauri and spatially resolved the Br $\gamma$  emission as well as nine transitions from the hydrogen Pfund series (Pf14-22). Comparing the spatial origin of the different line transitions, we find that the Brackett (Br $\gamma$ ), Pfund (Pf14-17), and Balmer (H $\alpha$ ) lines originate from different stellocentric radii ( $R_{\text{cont}} < R_{\text{Pf}} < R_{\text{Br}\gamma} \sim R_{\text{H}\alpha}$ ), which we can reproduce with an LTE line radiative transfer computation. Discussing different disk-formation scenarios, we conclude that our constraints are inconsistent with wind compression models predicting a strong outflowing velocity component, but support viscous decretion disk models, where the Keplerian-rotating disk is replenished with material from the near-critical rotating star.

*Subject headings:* circumstellar matter – stars: emission-line, Be – stars: individual ( $\beta$  CMi,  $\zeta$  Tau) – stars: fundamental parameters – techniques: interferometric

## 1. INTRODUCTION

Classical Be stars are main-sequence (or near main-sequence) B-type stars associated with hydrogen line emission, indicating the presence of ionized circumstellar gas, which is believed to be arranged in an equatorial disk-like structure. Optical/infrared spectroscopic observations have revealed non-radial pulsations (e.g. Rivinius et al. 1998), which might provide a way to feed material from the photosphere to the inner disk. Polarimetric studies (e.g. Draper et al. 2011) can constrain the disk density structure. Further unique insights into the structure and physics of these disks can also be obtained with interferometry at visual and infrared wavelengths, allowing one to unravel the inner disk structure on scales of a few stellar radii directly. For instance, interferometric studies have allowed associating quasi-cyclic variations in the ratio between the blue- and redshifted wing of the H $\alpha$ -line emission ( $V/R$  variability) with global oscillations in the circumstellar disk, likely in the form of a one-armed spiral density pattern (e.g. Vakili et al. 1998; Štefl et al. 2009; Carciofi et al. 2009;

Schaefer et al. 2010). Besides studies on the disk continuum geometry, interferometric observations in spectral lines have provided the first direct constraints on the gas kinematics, in particular for the hydrogen spectral lines of the Balmer (e.g. Quirrenbach et al. 1994; Vakili et al. 1998; Tycner et al. 2005; Delaa et al. 2011), Brackett (Meilland et al. 2007, 2011), and Pfund series (Pott et al. 2010). These studies have provided growing evidence that the disks around classical Be stars exhibit a near-Keplerian rotation profile, which might allow to effectively rule out several disk-formation scenarios (e.g. see review by Carciofi 2010). However, most earlier studies using spectro-interferometry were limited in terms of spectral resolution or baseline position angle (PA) coverage, leaving significant uncertainties about the detailed gas velocity field and the evidence to distinguish between a purely rotational versus an expanding velocity component in the disk. Obtaining such evidence is essential in order to decide between different scenarios which have been proposed to explain the disk formation mechanism, including radiatively driven winds, ram pressure or magnetically induced wind compression, and viscous decretion (Porter & Rivinius 2003). Furthermore, recently there has been a controversy about the appearance of a phase inversion in spectro-interferometric observations

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of several classical Be stars, which triggered speculations about secondary dynamical effects or the need for an additional kinematical component beyond the canonical star+disk paradigm (Steff et al. 2011).

Here, we present near-infrared spectro-interferometric observations on the classical Be star  $\beta$  CMi with a high spectral resolution of  $R = 12,000$  in the hydrogen Br $\gamma$ -line, enabling us to constrain the rotation profile directly. In addition, we observed the classical Be star  $\zeta$  Tau for the first time in multiple hydrogen line transitions (Br $\gamma$  and Pfund lines), providing direct information about the excitation structure within the disk.

In the following, we present our CHARA and VLTI interferometric observations (Sect. 2). The observations in spectral lines are then first interpreted using a model-independent photocenter analysis approach (Sect. 3). In Sect. 4, we present our continuum and kinematical modeling on  $\beta$  CMi, followed by our discussion of the results on  $\zeta$  Tau (Sect. 5). Finally, we summarize our findings in Sect. 6.

## 2. OBSERVATIONS

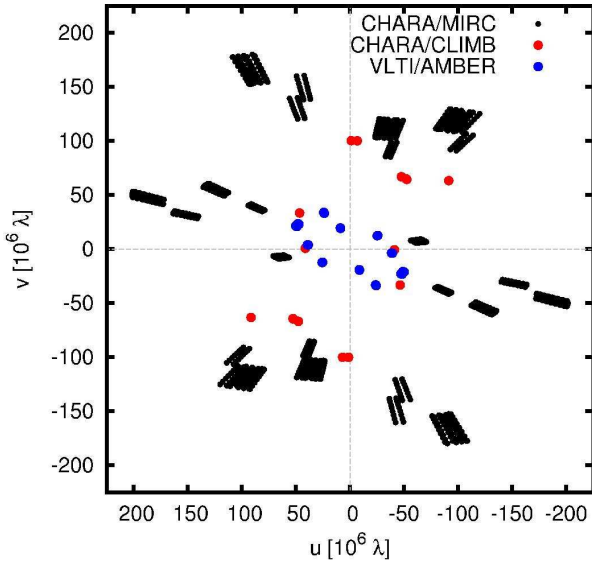


FIG. 1.—  $uv$ -coverage achieved with our CHARA/MIRC ( $H$ -band), CHARA/CLIMB ( $K$ -band), and VLTI/AMBER ( $K$ -band) interferometric observations on  $\beta$  CMi.

Our near-infrared  $H$ -band continuum observations on  $\beta$  CMi were obtained using the CHARA array (ten Brummelaar et al. 2005), which is operated by Georgia State University. The MIRC beam combiner (Monnier et al. 2006b) allowed us to combine the light from four of the six CHARA 1 m-telescopes simultaneously, yielding baseline lengths of up to 330 m, with good baseline coverage (Fig. 1). The MIRC data cover the  $H$ -band with low spectral dispersion ( $R = 35$ ) and was reduced using the University of Michigan MIRC data reduction pipeline (Monnier et al. 2007).

In order to investigate the  $K$ -band disk geometry of  $\beta$  CMi, we employed the CHARA/CLIMB 3-telescope beam combiner (Sturmann et al. 2010). Visibilities and closure phases were derived using the “redclimb” and “reduceir” software. Besides the statistical errors, we also add a calibration uncertainty of 0.05 for the derived

NOTE. — Column 5 (NP) denotes the number of pointings obtained on the science star. For the calibrators, we assume the following uniform disk (UD) diameters: For HD 25490 ( $0.536 \pm 0.037$  mas), HD 43042 ( $0.62 \pm 0.05$  mas), HD 71095 ( $2.009 \pm 0.139$  mas), and HD 73262 ( $0.47 \pm 0.03$  mas) we use the diameter estimates from searchCal (Bonneau et al. 2006). The UD diameters of HD 79469 ( $0.471 \pm 0.027$  mas) and HD 97633 ( $0.68 \pm 0.06$  mas) have been estimated by averaging the estimates from three independent photometric methods (Barnes et al. 1978; Bonneau et al. 2006; Kervella & Fouqué 2008).

TABLE 1  
OBSERVATION LOG OF OUR CHARA AND VLTI OBSERVATIONS.

Target	Date (UT)	Instrument	Spectral mode	NP	DIT [s]	Telescope configuration	Calibrator(s)
$\beta$ CMi	2009-12-31	AMBER	HR-K 2.172	1	1	UT1-UT2-UT4	HD 71095
	2010-04-23	AMBER	HR-K 2.172	1	6	K0-G1-A0	HD 71095
	2008-12-02	MIRC	H35	1	-	S1-E1-W1-W2	HD 25490, HD 79469, HD 97633
	2009-11-10	MIRC	H35	3	-	S1-E1-W1-W2	HD 43042
	2010-11-02	MIRC	H35	3	-	S1-E1-W1-W2	HD 43042
	2010-11-03	MIRC	H35	2	-	S2-E2-W1-W1	HD 43042
	2010-12-14	MIRC	H35	1	-	S1-E2-W1-W2	HD 79469, HD 97633
	2010-11-30	CLIMB	K	2	-	S2-W1-W2	HD 73262
	2010-12-02	CLIMB	K	1	-	S1-E2-W1	HD 73262
	2010-12-03	CLIMB	K	4	-	S2-E2-W2	HD 73262
$\zeta$ Tau	2010-01-01	AMBER	MR-K 2.3	1	0.2	UT1-UT2-UT4	HD 71095

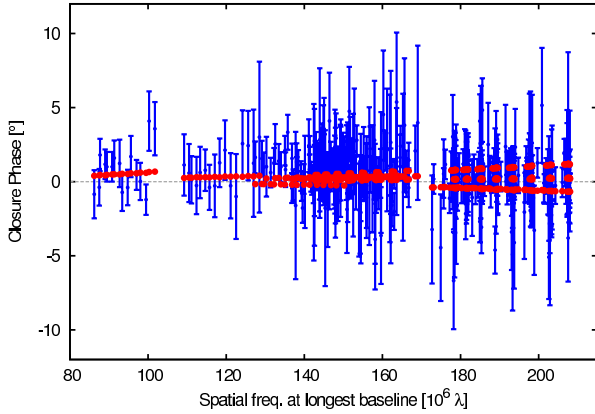


FIG. 2.— Closure phases measured with the MIRC beam combiner on  $\beta$  CMi (blue data points), overplotted with the model predictions from our best-fit  $H$ -band model (red data points; Sect. 4.2).

visibilities, which represents an empirical value for the typical scatter in the instrument transfer function.

Spectro-interferometric observations with medium (MR mode,  $R = 1500$ ) and high spectral dispersion (HR mode,  $R = 12000$ ) were obtained with the *Very Large Telescope Interferometer* (VLTI) of the European Southern Observatory and the AMBER 3-telescope beam combiner instrument (Petrov et al. 2007). The AMBER observations on  $\beta$  CMi were recorded using three 8.2 m unit telescopes (2009-12-31) and three 1.8 m auxiliary telescopes (2010-04-23), respectively. The atmospheric piston was stabilized using the FINITO fringe tracker (Le Bouquin et al. 2008), which allowed us to use long detector integration times (DITs) of 1 s and 6 s and to record data with high spectral dispersion around the hydrogen Br $\gamma$ -line ( $\lambda_{\text{Br}\gamma}^{\text{vacuum}} = 2.166078 \mu\text{m}$ ). Due to the presence of residual phase jitter, the absolute visibility calibration of the AMBER data is not reliable, while the important wavelength-differential observables are not affected. Spectra and wavelength-differential visibilities and phases (DPs, Fig. 3) were derived from the AMBER data using the amdlib (V3.0) data reduction software (Tatulli et al. 2007; Chelli et al. 2009). The wavelength calibration was done using atmospheric telluric features close to the Br $\gamma$ -line (yielding an accuracy of approximately 1 spectral channel) and by applying a heliocentric-barycentric system correction using heliocentric velocities of  $+5.97 \text{ km s}^{-1}$  (2009-12-31) and  $-28.61 \text{ km s}^{-1}$  (2010-04-23), respectively. For the systemic velocity we assume  $+22.0 \text{ km s}^{-1}$  (Duflot et al. 1995).

From all CHARA and VLTI interferometric observations, we also derived closure phases (CPs). Both the  $H$ - and  $K$ -band continuum closure phases are consistent with zero on a  $2\sigma$ -level, which leads us to conclude that the brightness distribution does not show significant indications for deviations from centro-symmetry. The most constraining CPs have been recorded with MIRC in the  $H$ -band, which are shown in Fig. 2. Our AMBER HR measurements from 2009-12-31 provide us also with a CP measurement in the Br $\gamma$ -line of  $\beta$  CMi, while the derived CPs from the 2010-04-23 dataset are rather noisy and are therefore not included for our model fits.

$\zeta$  Tau was observed on 2010-01-01 using AMBER's

MR-mode covering the upper  $K$ -band ( $2.12$  to  $2.46 \mu\text{m}$ ). The assumed systemic velocity for  $\zeta$  Tau is  $+21.8 \text{ km s}^{-1}$  (Duflot et al. 1995). In the  $\zeta$  Tau data, we detected not only the Br $\gamma$  transition (7-4), but also hydrogen Pfund transitions (Fig. 4, 1st row), including clear detection of Pf14 (19-5,  $2.4477 \mu\text{m}$ ) to Pf22 (27-5,  $2.3591 \mu\text{m}$ ). Higher Pfund transitions are also present in the spectrum, but cannot be clearly separated due to the wide, double-peaked profile of the individual lines. For  $\zeta$  Tau, the derived CPs are too noisy to provide additional information and are therefore not included in our further analysis.

Details about the observational setup for all interferometric observations are listed in Tab. 1. Each observation on a science star was accompanied by observations on interferometric calibrators, allowing us to monitor and correct for the atmospheric and instrumental transfer function.

### 3. PHOTOCENTER ANALYSIS

Differential phases provide unique information about small-scale (sub-mas) photocenter displacements between the blue- and red-shifted line wings. These displacements provide a very sensitive measure of the gas kinematics on scales of a few stellar radii. In a first analysis step, we reconstruct the on-sky 2-D photocenter displacement from the measured DPs by solving the system of linear equations

$$\vec{p} = -\frac{\phi_i}{2\pi} \cdot \frac{\lambda}{\vec{B}_i}, \quad (1)$$

where  $\phi_i$  is the differential phase measured on baseline  $i$ ,  $\vec{B}_i$  is the corresponding baseline vector, and  $\lambda$  is the central wavelength (Le Bouquin et al. 2009). The derived photocenter plots for  $\beta$  CMi and  $\zeta$  Tau are shown in Figs. 3 and 4 (*middle panel*), respectively, and clearly reveal rotation-dominated velocity fields for both objects, as indicated by the linear alignment of the photocenter vectors for the different gas velocities and the opposite sign of the photocenter displacement for the blue- and red-shifted emission. In order to associate the DP with the on-sky orientation, it is necessary to calibrate the sign of the DP measurements. For this purpose, we reprocessed the  $\zeta$  Tau VLTI/AMBER data set presented by Štefl et al. (2009) and calibrated our DP sign in order to match the published on-sky orientation.

For  $\beta$  CMi, we determine the position angles  $\theta$  for the disk rotation plane at the two epochs to be  $138.3 \pm 1.5^\circ$  (2009-12-31) and  $141.7 \pm 1.4^\circ$  (2010-04-22). The photocenter vectors corresponding to our highest-SNR observation (2009-12-31) show an interesting arc-like structure in the red-shifted line wing (Fig. 3, *bottom left*), where the photocenter vectors corresponding to low gas velocities are above the derived disk plane, while the high velocities are displaced in the opposite direction. Although the significance of this pattern is still only marginal in our data, we speculate that this pattern might result from opacity effects, with the more distant parts of the disk appearing fainter than the disk parts facing the observer. Such an obscuration screen would displace the photocenter perpendicular to the disk plane, where the amplitude of the displacement is stronger for lower gas velocities, since the low-velocity emission is distributed

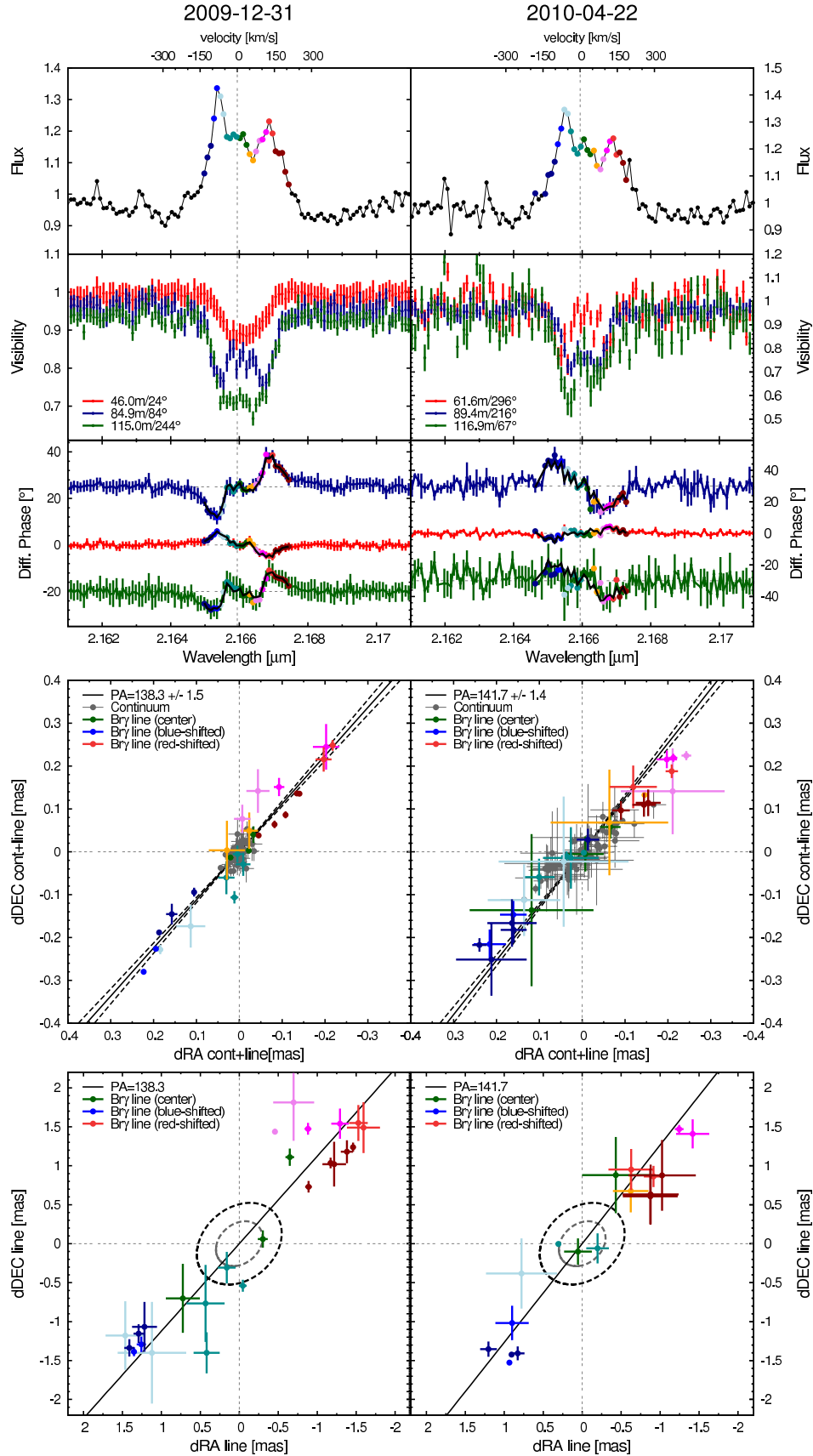


FIG. 3.— *Upper panel:* VLTI/AMBER spectra (1st row), visibilities (2nd row), and DPs (3rd row) measured on  $\beta$  CMi for the epochs 2009-12-31 (left) and 2010-04-22 (right). *Middle panel:* From the measured DPs, we derive for each spectral channel the 2-D photocenter displacement vector (including continuum and line contributions; East is plotted left and North is up). The flux and DP data points as well as the derived photocenter vectors have been color-coded based on the Doppler velocity (see the upper panel to relate each color to a wavelength). The DPs corresponding to the astrometric solutions are shown in the middle panel as solid black lines. *Bottom panel:* Using the procedure outlined in Sect. 3, we corrected for the continuum contributions in the line spectral channels, revealing the photocenter displacement corresponding to the line emission only. Besides the determined photocenter offsets, for comparison we also show the size of the  $H$ -band continuum-emitting disk (black ellipse, FWHM Gaussian, Sect. 4.2) and of the stellar photosphere (grey ellipse, Sect. 4.1).

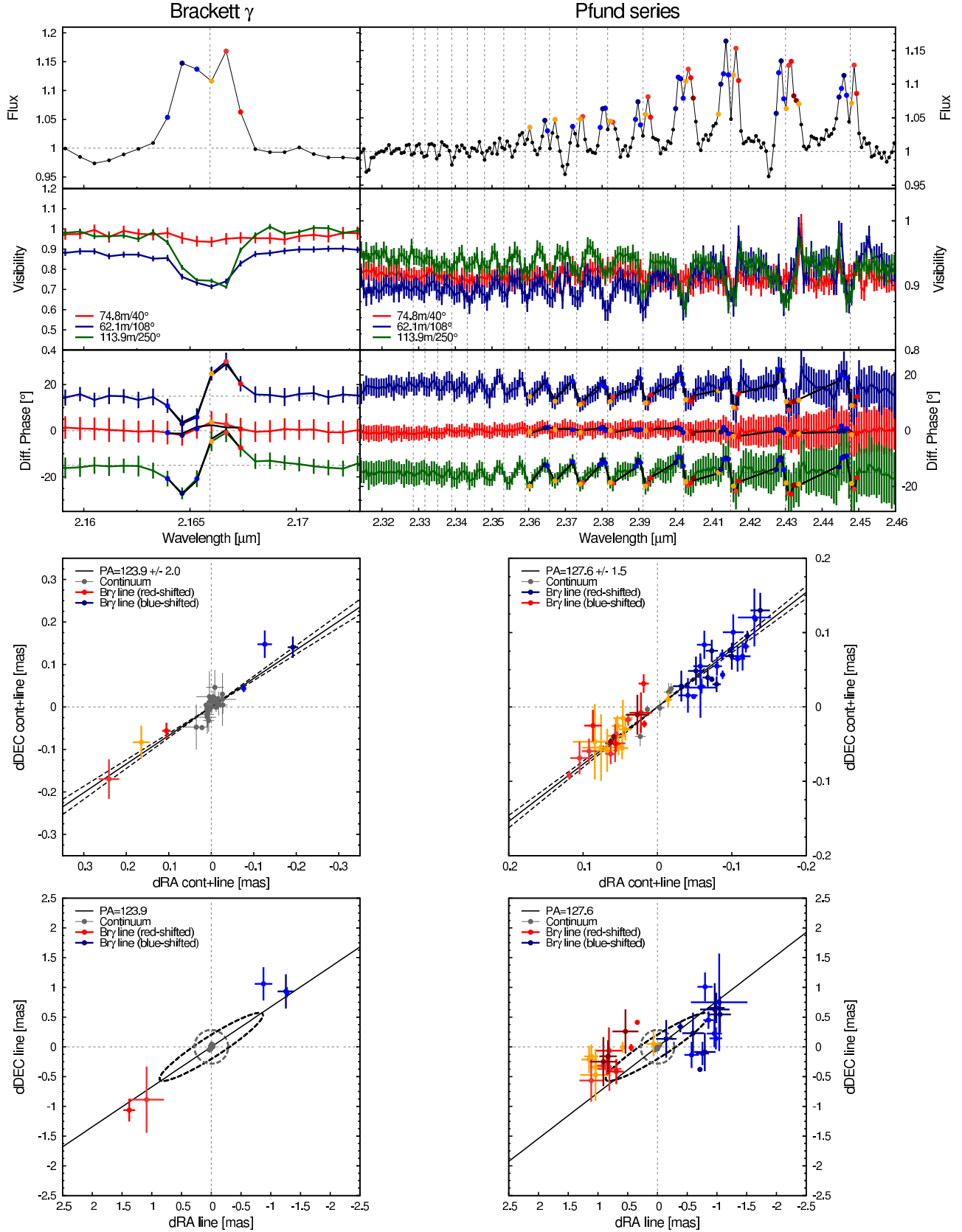


FIG. 4.— Upper panel: VLTI/AMBER spectra (1st row), visibilities (2nd row), and DPs (3rd row) measured on  $\zeta$  Tau for the epoch 2010-01-01 in the Br $\gamma$  (left) and Pf14-22 transitions (right). Middle/bottom panel: Photometer displacement vectors derived from the measured DPs (including line and continuum emission, middle panel) and the continuum-corrected DPs (tracing the line emission only, bottom panel). For comparison, we also show the size of the  $H$ -band continuum-emitting disk, as determined by Schaefer et al. (2010) for epoch 2009-11-10 (black ellipse, FWHM Gaussian) and of the stellar photosphere (grey ellipse).



over a more extended region. Accordingly, the displacement would be strongest at zero velocities, and then symmetrically decrease towards higher velocities. The superposition of this weak displacement (perpendicular to the disk plane) with the displacement due to Keplerian rotation (parallel to the disk plane) might result in the observed arc-shaped structure, which would also provide a unique tool to determine the orientation of the disk in space and the disk rotation sense. In the case of  $\beta$  CMi, this implies that the north-eastern part of the disk is facing towards the observer (based on the displacement of the low-velocity channels in this direction) and that the disk is in clockwise rotation (based on the location of the red-shifted photocenter displacements in the north-western quadrant).

Another intriguing feature in the  $\beta$  CMi data is the  $W$ -shaped profile, which we observe in the wavelength-dependent visibilities and DP on our longest interferometric baselines (Fig. 3, *2nd* and *3rd* row). Likely, this profile indicates that the visibility function of the line-emitting region passes through a visibility null and transits from the first to the second visibility lobe. Since this effect would reverse the direction of the photocenter vector, we manually correct the phase sign in these corresponding spectral channels close to the line center. For  $\zeta$  Tau, we can determine the rotation axis for the Br $\gamma$  ( $123.9 \pm 2.0^\circ$ ) and the nine Pf14-Pf22 transitions ( $127.6 \pm 1.5^\circ$ ) separately and find that the line-emitting gas rotates in the same disk plane within the observational uncertainties of  $\sim 2^\circ$ .

Using the aforementioned procedure, it is possible to reliably measure the direction of the photocenter displacement, while the length of the displacement vector is biased by the contributions from the underlying continuum emission. In order to remove these contributions from the measured observables ( $F$ ,  $V$ ,  $\phi$ ), we apply the method outlined by Weigelt et al. (2007) and interpolate the continuum flux ( $F_c$ ) and continuum visibility ( $F_c$ ) from the adjacent continuum. The visibility and DP of the pure line emitting-region ( $V_l$ ,  $\phi_l$ ) are then given by

$$|F_l V_l|^2 = |F V|^2 + |F_c V_c|^2 - 2 \cdot F V \cdot F_c V_c \cdot \cos \phi \quad (2)$$

$$\sin \phi_l = \sin \phi \frac{|F V|}{|F_l V_l|}, \quad (3)$$

where  $F_l = F - F_c$  denotes the flux contribution from the spectral line. The continuum-corrected DPs are then used to derive the photocenter displacement of the pure line-emitting region (Figs. 3 and 4, *bottom panel*). Applying this correction will provide the real centroid offset of the line emission, but also introduce noise from the visibility and flux measurements, resulting in an increased scatter in the position angle distribution. Therefore, we decided to measure the position angle of the rotation axis from the uncorrected line+continuum photocenter displacements (Figs. 3 and 4, *middle panel*), while the continuum-corrected photocenter displacements (*bottom panel*) will be used later on to construct a position-velocity diagram and to compare the spatial origin in different line tracers.

#### 4. DISCUSSION ON $\beta$ CANIS MINORIS

$\beta$  CMi (HR 2845) is a relatively quiet B8V-type classical Be star located at a distance of  $52.2^{+2.4}_{-2.2}$  pc

(Tycner et al. 2005). Ground-based photometric monitoring has provided no clear indications for significant variability (Pavlovski et al. 1997), while there is some marginal evidence for long-term variations in the H $\alpha$ -profile (Pollmann 2002; Hesselbach 2009).

In the following, we present a refined model for the photospheric emission of  $\beta$  CMi taking the near-critical stellar rotation into account (Sect. 4.1), followed by our modeling of the interferometric data in continuum emission (Sect. 4.2). Given that the disk rotation signatures in the Br $\gamma$  emission line are overlayed on the rotation signatures of the star (in Br $\gamma$  absorption), we investigate the influence of the stellar rotation on our measurements (Sect. 4.3). In the following, we construct a position-velocity diagram (Sect. 4.4) and present a full kinematical modeling in Sect. 4.5.

##### 4.1. Constraining the stellar parameters

In order to obtain a model for the photospheric emission of  $\beta$  CMi, we searched for a consistent set of stellar parameters, taking the evidence for near-critical rotation into account (Saio et al. 2007). For this purpose, we employed our rapid rotator code (Monnier et al. 2007; Che et al. 2011), which simulates the stellar oblateness and surface temperature distribution using the modified von Zeipel theorem (gravity darkening coefficient = 0.188), and computes model images in the continuum and in the photospheric Br $\gamma$  absorption line.

In order to constrain the stellar parameters, we computed a small parameter grid, in which we systematically varied the inclination angle  $i$  and the fractional angular velocity  $\omega/\omega_{\text{crit}}$  (where  $\omega_{\text{crit}}$  is the critical angular velocity). For a given inclination angle and fractional angular velocity, we start our iterative process by assuming a stellar mass  $M_*$ , from which we compute the polar radius  $R_{\text{pole}}$  using the rotation velocity  $v \sin i = 244 \pm 6 \text{ km s}^{-1}$  (Yudin 2001). The polar temperature  $T_{\text{pole}}$  is derived from the  $V$ -band magnitude ( $V = 2.89$ ). Any disk continuum flux in the optical band will tend to dilute the photospheric spectrum causing absorption lines to appear weaker than expected. We checked for such line reduction in blue spectra of

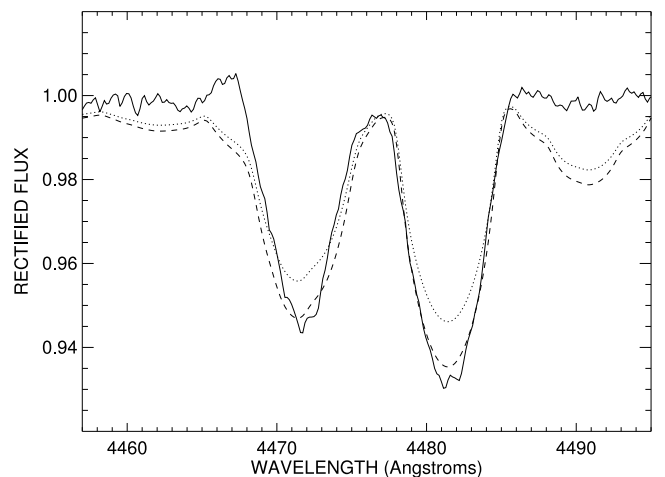


FIG. 5.— Spectrum of the He I 447.1 nm and Mg II 448.1 nm line of  $\beta$  CMi (solid line). The dashed line shows a photospheric model without disk contributions, while the dotted line corresponds to a model where 20% disk continuum contributions have been added.

TABLE 2  
GRID OF STELLAR PARAMETERS, OBTAINED FOR  $\beta$  CMi USING OUR FAST ROTATOR MODEL (SECT. 4.1).

Model	$i$ [ $^\circ$ ]	$\omega/\omega_{\text{crit}}$	$M_\star$ [ $M_\odot$ ]	$R_{\text{pole}}$ [ $R_\odot$ ]	$R_{\text{eq}}$ [ $R_\odot$ ]	$T_{\text{pole}}$ [K]	$P$ [days]	$L^{\text{app}}$ [ $L_\odot$ ]	$T_{\text{eff}}^{\text{app}}$ [K]
A	45	0.99	3.70	3.32	4.61	13200	0.676	267	11558
B	45	0.96	4.11	3.03	3.94	14600	0.577	343	13085
C	45	0.93	4.47	2.87	3.59	15650	0.527	413	14235
D	40	0.99	3.98	2.95	4.10	14550	0.546	337	12871
E	40	0.96	4.49	2.74	3.56	16050	0.475	435	14505
F	40	0.93	4.91	2.60	3.25	17300	0.433	532	15847
G	35	0.99	4.43	2.62	3.64	16200	0.433	437	14447
H	35	0.96	5.03	2.44	3.17	18100	0.377	590	16474

$\beta$  CMi obtained by Grundstrom (2007). Fig. 5 shows the observed profiles (*solid line*) of He I 447.1 nm and Mg II 448.1 nm based upon the average of 11 spectra made between 2005 and 2009 with the Kitt Peak National Observatory Coude Feed Telescope (spectral resolution  $R = 12,500$ ). The observed spectrum is compared with a theoretical flux spectrum (*dashed line*) from the grid of LTE models by Rodríguez-Merino et al. (2005). We adopted average parameters for the visible hemisphere of the star of  $T_{\text{eff}} = 11,800$  K,  $\log g = 3.8$ , and  $V \sin i = 230 \text{ km s}^{-1}$  (Frémat et al. 2005), and the model spectrum was convolved with a simple rotational broadening function for a linear limb darkening coefficient of  $\epsilon = 0.42$  (Wade & Rucinski 1985). Note that this approach assumes a spherical star, ignores gravity darkening, and neglects changes in the local intensity spectrum with orientation, but these simplifications are acceptable for our purpose of checking for systematic line depth differences. We see that the match of the Mg II  $\lambda 4481$  profile is satisfactory, but the wings of the He I  $\lambda 4471$  line appear too high, probably due to low level emission in this transition. Adding a 20% disk flux contribution and renormalizing the spectrum yields a diluted version (*dot-dashed line*), which appears to be much too weak compared

with the observed spectrum. This comparison suggests that the disk continuum emission in the blue range is negligible, and thus we will ignore any disk contribution to the optical flux in the following discussion.

With the derived polar radius and bolometric luminosity, we are able to locate the stellar position on the HR diagram, after correcting for the rotational effect (Che et al. 2011). The stellar mass from the HR diagram is used as initial value for the next iteration step, until convergence between the assumed mass and the mass estimated from the HR diagram is reached.

For each parameter combination, our model provides the rotation period, apparent luminosity  $L^{\text{app}}$ , and apparent effective temperature  $T_{\text{eff}}^{\text{app}}$ , as listed in Tab. 2. Comparing the model effective temperature and luminosity with the observational constraints for  $\beta$  CMi ( $T_{\text{eff}} = 12,050$  K;  $L = 195 \pm 60 L_\odot$ ; Saio et al. 2007), we find good agreement for models A, B, and D. Due to its consistency with the inclination and effective temperature (Sects. 4.2 and 4.5), we favor model D, suggesting  $i = 40^\circ$ ,  $\omega/\omega_{\text{crit}} = 0.99$ ,  $T_{\text{pole}} = 14,550$ ,  $v \sin i = 244 \text{ km s}^{-1}$ ,  $T_{\text{eff}}^{\text{app}} = 12,871$  K, and a polar and equatorial radius of  $R_{\text{pole}} = 2.95 R_\odot = 0.26 \text{ mas}$  and  $R_{\text{eq}} = 4.10 R_\odot = 0.36 \text{ mas}$ , respectively. We plot the values corresponding to our model grid in the HR diagram shown in Fig. 6 and compare it to the evolutionary tracks from Yi et al. (2003) and Demarque et al. (2004), yielding an evolutionary age of  $\sim 0.1 - 0.15$  Gyr.

Obviously, an important input parameter for our modeling procedure is the rotation velocity  $v \sin i$ , for which we use an average value from the literature. Townsend et al. (2004) argued that most  $v \sin i$  measurements on Be stars might systematically underestimate the true projected rotation value due to the effect of gravity darkening. In order to test this scenario, we have artificially increased the measured  $v \sin i$ -value by 10% and repeated our modeling procedure. We find that the increase in  $v \sin i$  results in a significant increase in the stellar mass, apparent effective temperature, and apparent luminosity, making the model prediction much less consistent with the observed values. Therefore, we suggest that in the case of  $\beta$  CMi, the  $v \sin i$ -measurements are not significantly biased by gravity darkening, which is likely a result of the intermediate inclination angle of  $\beta$  CMi (the effect would be stronger for an equator-on viewing angle) and the use of our moderate gravity darkening law ( $T_{\text{eff}} \propto g_{\text{eff}}^{0.188}$ , Che et al. 2011), which is likely more relativistic than the classical von Zeipel darkening coefficient ( $T_{\text{eff}} \propto g_{\text{eff}}^{0.25}$ ) adopted by Townsend et al.

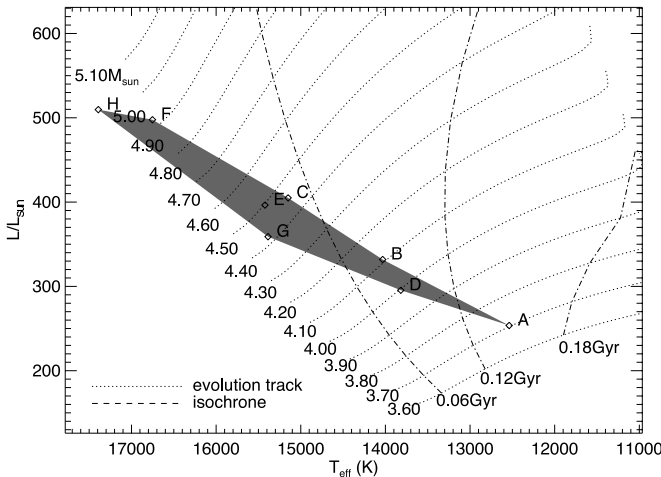


FIG. 6.— Hertzsprung-Russell diagram, including the solutions from our small rapid rotator grid for  $\beta$  CMi (Tab. 2). The numbers on the left side give the stellar mass of the corresponding evolution track, while the numbers at the bottom give the evolutionary age corresponding to the shown isochrones Yi et al. (2003); Demarque et al. (2004). As outlined in Sect. 4.1, only models A, B, and D satisfy the observational constraints on the stellar effective temperature. Furthermore, model D satisfies best our inclination obtained with interferometry ( $38.5 \pm 1^\circ$ ).





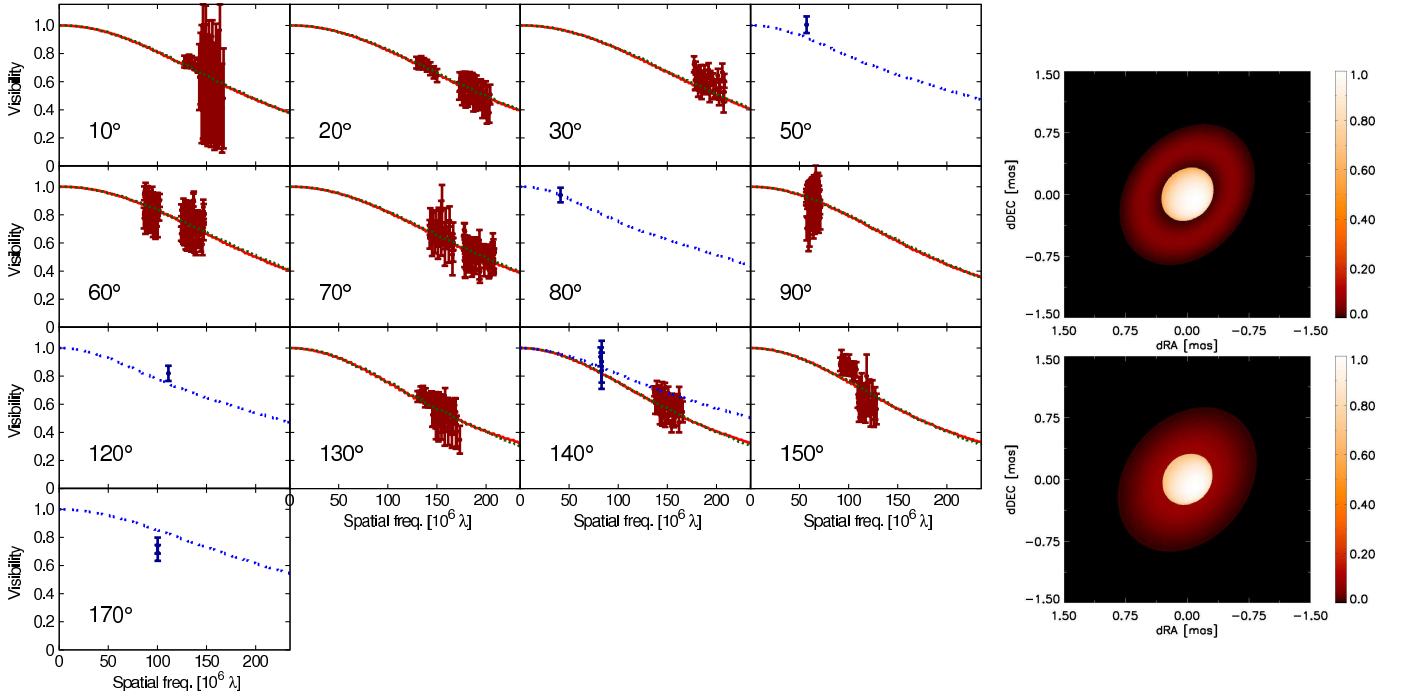


FIG. 7.— *Left*: Continuum visibilities measured on  $\beta$  CMi with CHARA/MIRC in the  $H$ -band (red points) and CHARA/CLIMB in the  $K$ -band (blue points). In order to show the position-angle dependence of the visibility function, we have binned the data by the projected baseline PA (number in lower left corner of each panel). For the  $H$ -band, the best-fit ring model (red line) and Gaussian model (green dotted line) are shown, with both models yielding nearly indistinguishable visibility profiles. For  $K$ -band, we show the best-fit Gaussian model (blue dashed line; see Sect. 4.2). *Right*: Intensity distribution corresponding to our best-fit ring (*top*) and Gaussian model (*bottom*).

sure phases ( $\lesssim 1.2^\circ$ ), consistent with the measurement (Fig. 2).

The determined disk position angles of  $139.2^\circ$  (ring model) and  $139.3^\circ$  (Gaussian model) are in excellent agreement with the gas disk rotation axis determined with our VLTI/AMBER photocenter analysis ( $140.0 \pm 1.7^\circ$ ).

Using the detailed information about the  $H$ -band geometry obtained with our extensive MIRC data set, we then fitted our Gaussian model to the CLIMB data to determine the  $K$ -band continuum geometry. Given the lower amount of observational constraints for the  $K$ -band, we treated only the disk-to-star flux ratio  $(F_{\text{disk}}/F_\star)_K$  as a free parameters and kept the remaining parameters fixed. The resulting best-fit visibility curves are shown in Fig. 7. With  $(F_{\text{disk}}/F_\star)_K = 0.25^{+0.09}_{-0.08}$ , we do not find any significant deviations between the  $K$ -band and  $H$ -band flux ratio.

#### 4.3. Investigating the differential phase signatures of the stellar rotation in the $\text{Br}\gamma$ absorption line

With the currently achievable accuracy, DP measurements can already reveal photocenter displacements two to three orders smaller than the formal angular resolution ( $\lambda/2B$ ). For instance, our 2009-12-31 AMBER observations exhibit a DP accuracy of  $\sim 0.8''$  (standard deviation over all continuum spectral channels), corresponding to a photocenter displacement of  $\sim 8$  micro-arcsecond ( $0.008$  mas) on the employed 115m baseline. Given that this is much smaller than the equatorial stellar diameter ( $\sim 660$  micro-arcsecond), it is important to investigate whether the measured DP signatures might also contain contributions from the photospheric  $\text{Br}\gamma$  absorption line,

which is tracing the stellar rotation and is underlying the  $\text{Br}\gamma$ -line emission.

In order to simulate the effect of stellar rotation on our interferometric observables, we employ our rapid rotator code and compute the stellar surface brightness distribution around the  $\text{Br}\gamma$ -line with a similar resolution as our AMBER observations (Fig. 8, *top*). For this, we align the stellar rotation axis with the measured rotation axis ( $\theta = 227.5^\circ$ , Sect. 3).  $\text{Br}\gamma$  absorption on one side of the photosphere will shift the photocenter towards the opposite direction. Taking this into account, we adjust the orientation of our model photosphere, so that the  $\text{Br}\gamma$  absorption in the blue-shifted line wing matches the measured direction of the photocenter displacement at red-shifted wavelengths. In addition to the photospheric emission, we include disk emission in our model, assuming  $(F_{\text{disk}}/F_\star)_K = 0.25$ , as determined in Sect. 4.2.

From the model, we compute the corresponding DPs and find signatures  $\phi \lesssim 0.5^\circ$  on all baselines (Fig. 8, *bottom*). This result reflects the fact that the fraction of the photospheric absorption to the total  $K$ -band emission is rather small ( $< 8\%$ ) and that the absorption appears as a rather diffuse structure on the extended stellar surface. We conclude that the photocenter signatures due to stellar rotation (in photospheric absorption lines) are a minor effect compared to the rotation signatures of the disk (in line emission), but nevertheless result in DP signatures comparable to the currently achievable DP accuracy and are therefore important, in particular for more detailed future studies.

#### 4.4. Constraints on the disk kinematics from the position-velocity diagram

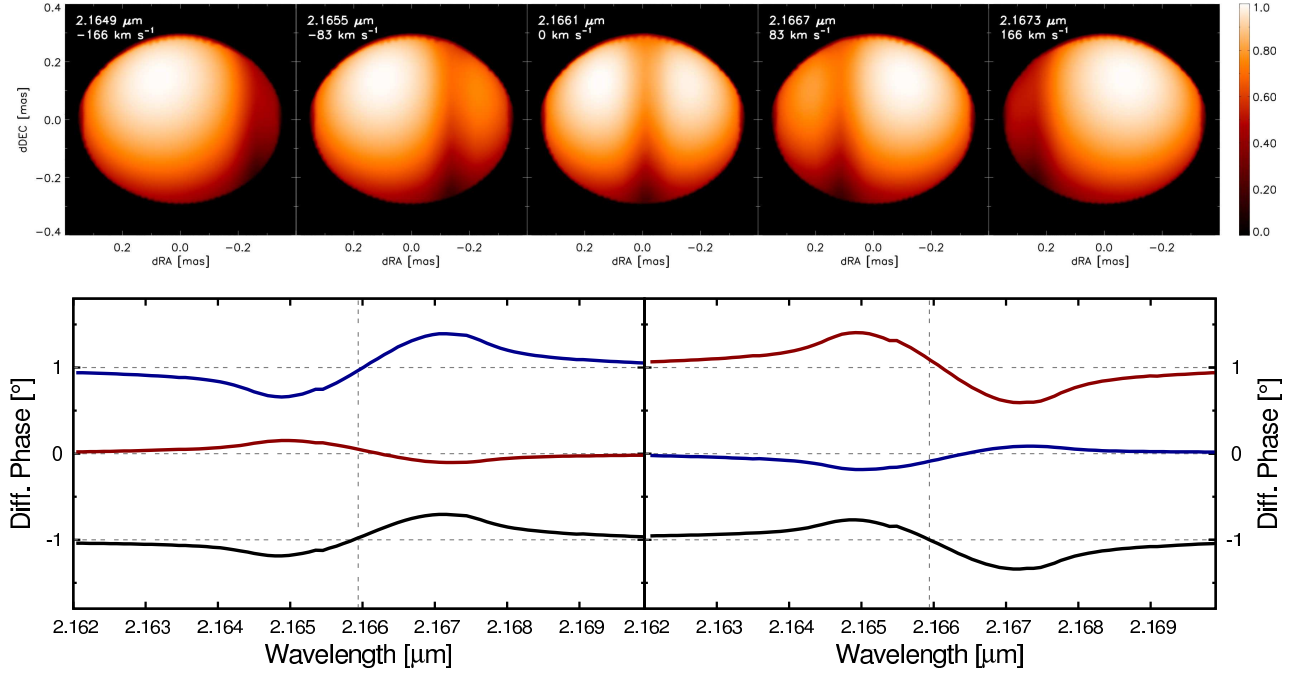


FIG. 8.— *Top*: Model images from our rapid rotator model for  $\beta$  CMi for some representative wavelengths (Sect. 4.3). *Bottom*: Model differential phases (including the stellar photosphere only) computed for the baselines employed by our  $\beta$  CMi VLTI measurements from 2009-12-31 (*left*) and 2010-04-23 (*right*).

Position-velocity ( $p-v$ ) diagrams provide a powerful tool for the interpretation of a disk velocity field and are commonly employed in radio astronomy to derive the rotation profile of circumnuclear galactic disks or protostellar disks using, for instance, maser (e.g. Miyoshi et al. 1995; Pestalozzi et al. 2009) or molecular emission tracers (e.g. Sofue & Rubin 2001; Isella et al. 2007). Using our continuum-corrected photocenter displacement measurements, we can construct an equivalent diagram from our VLTI interferometric data by measuring the length of the continuum-corrected photocenter displacement vectors projected on the disk plane  $\theta = 140.0^\circ$ . Performing this projection on the disk plane also allows us to avoid a potential bias due to opacity effects, since these effects would move the photocenter only perpendicular to the disk plane (Sect. 3). The resulting  $p-v$  diagram (Fig. 9) shows a symmetric rotation curve, which we interpret using the simple model of a thin Keplerian-rotating disk (Weintroub et al. 2008). The disk extends from an inner ( $R_{\text{in}}$ ) to an outer ( $R_{\text{out}}$ ) radius and the line-of-sight (LOS) velocity  $v$  of the emission element located at radius  $r$  and at angle  $\vartheta$  in the disk plane is given by

$$v_{\text{kep}}(r, \vartheta) = \sqrt{\frac{GM_\star}{r}} \sin \vartheta \cdot \cos i, \quad (4)$$

where  $\vartheta$  is measured against the LOS and  $G$  is the gravitational constant. Considering only the emission from the outer-most disk annulus (i.e. at  $R_{\text{out}}$  and  $\vartheta = -\pi \dots +\pi$ ) will result in a straight line in the  $p-v$  diagram (line  $A-B$  in Fig. 9). Adding the emission from the remaining disk annuli ( $R_{\text{in}}$  to  $R_{\text{out}}$ ) will result in a characteristic “bowtie”-shaped filled region in the  $p-v$  diagram, such as commonly observed in CO imaging observations (Sofue & Rubin 2001; Isella et al. 2007). However, astrometric observations (such as radio masers or photocenter

displacements) trace the light barycenter of each annulus, corresponding to the  $B-C$  and  $A-D$  curves in Fig. 9 (for more details see Weintroub et al. 2008). We fit this simple model to the  $\beta$  CMi  $p-v$  diagram assuming a fixed inclination of  $i = 40^\circ$  (as determined in Sect. 4.2), which yields best agreement for  $M_\star = 3.6 \pm 0.3 M_\odot$ ,  $R_{\text{out}} = 2.4 \pm 0.2$  mas, and  $R_{\text{in}} < 1.0$  mas (solid curve in Fig. 9).

#### 4.5. Detailed modeling of the disk kinematics

By reducing the information content to purely astrometric data, the  $p-v$  diagram analysis method presented in the last section provides a very intuitive and powerful method to constrain the gas velocity field in the observed line tracer. In this section, we will use a more sophisticated model in order to make full use of the rich information contained in our spectro-interferometric observations, including spectra, wavelength-differential visibilities, DPs, and CPs. The Br $\gamma$ -emitting gas is assumed to be located in a thin disk plane, which is justified due to the expected small opening angle of Be star disks (e.g. Bjorkman & Cassinelli 1993) and the intermediate inclination angle under which  $\beta$  CMi is observed (Sect. 4.2).

The gas kinematics is parameterized with a rotation profile  $|\vec{v}(r)| = f_{\text{kep}}(R_{\text{ref}}) \cdot (r/R_{\text{ref}})^\beta$ , where  $\beta = -1$  for rotation with constant angular momentum,  $\beta = -0.5$  for Keplerian rotation,  $\beta = 0$  for constant rotation, and  $\beta = +1$  for solid body rotation (e.g. Stee 1996).  $f_{\text{kep}} = |\vec{v}(R_{\text{ref}})|/|\vec{v}_{\text{kep}}(R_{\text{ref}})|$  is the orbital velocity at the reference radius  $R_{\text{ref}} = 1$  AU, expressed as fraction of the Keplerian velocity  $|\vec{v}_{\text{kep}}(r)| = (GM_\star/r)^{-1/2}$ . In the model, the line emission extends from the stellar surface to an outer radius  $R_{\text{out}}$  with a radial power-law intensity profile  $I_l(r) \propto r^q$ . To include thermal line broadening in our kinematic model, we adopt a constant radial gas temperature of  $0.6 \cdot T_{\text{eff}}$  ( $=7860$  K for  $\beta$  CMi),

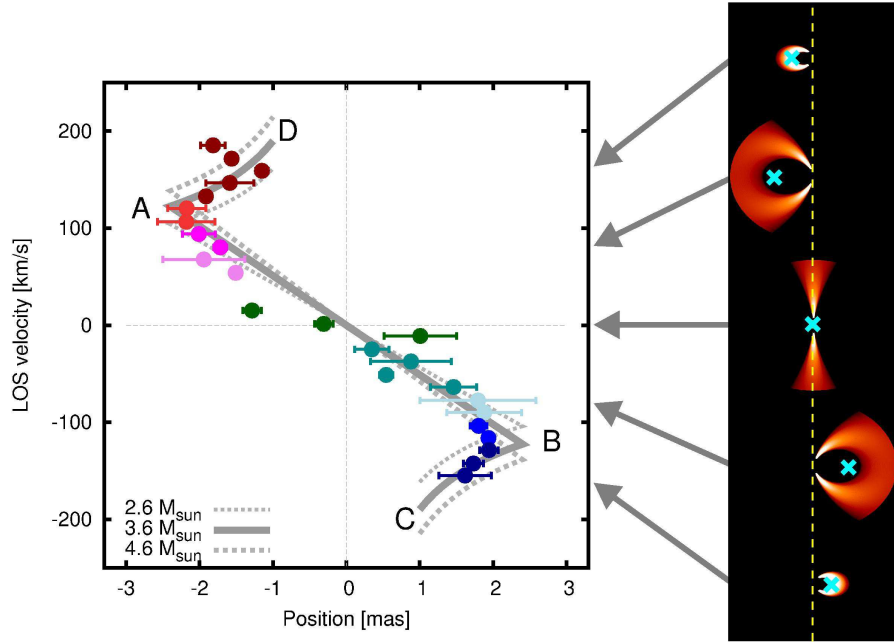


FIG. 9.— *Left*: Position-velocity diagram, derived from the continuum-corrected photocenter displacement vectors in Fig. 3 (*bottom, left*). The color of the data points matches the color-coding of the wavelength channels in Fig. 3. The model curves show the Keplerian rotation profile for three stellar masses with  $R_{\text{in}} = 0.68$  mas, and  $R_{\text{out}} = 2.4$  mas (Sect. 4.4). *Right*: In order to illustrate the interpretation of the  $p-v$  diagram, we show the Br $\gamma$ -line model images from our best-fit kinematical model (Sect. 4.5, Fig. 10) for some representative velocities and mark the corresponding photocenter displacement (corresponding to the position offset plotted on the abscissa in the diagram) in each channel map with a blue cross. The yellow vertical line marks the position of the star, which corresponds to the photocenter of the continuum emission and the zero position in the  $p-v$  diagram.

as suggested by the radiative transfer modeling from Carciofi & Bjorkman (2006).

We include both line and continuum emission in our model and compute the interferometric observables for our given VLTI array configurations and the covered wavelength channels. The wavelength-dependent model visibilities and phases are then fitted to the VLTI interferometric data using a reduced  $\chi^2_r$  goodness-of-fit estimator (see Kraus et al. 2009 for a definition), including our flux, visibility, DP, and CP constraints.

For the continuum emission, we assume the geometry determined with our CHARA observations (Sect. 4.2) and use this model to renormalize the AMBER continuum visibilities. In order to incorporate the underlying photospheric Br $\gamma$  absorption, we include the photosphere model discussed in Sect. 4.3, although, as discussed above, the influence on the differential phase ( $\lesssim 0.5^\circ$  on all baselines) is still within the measurement uncertainties.

The disk position angle is fixed to  $\theta = 140.0^\circ$ , as determined by our model-independent photocenter analysis (Sect. 3). The remaining six parameters in our modeling are the outer disk radius  $R_{\text{out}}$ , the inclination  $i$ , the stellar mass  $M_\star$ , the radial intensity power-law index  $q$ , the disk rotation index  $\beta$ , and the velocity at the reference radius  $R_{\text{ref}}$  (expressed as fraction of the Keplerian velocity,  $f_{\text{kep}}$ ). In a first step, we fix the velocity profile to Keplerian rotation ( $\beta = -0.5$ ,  $f_{\text{kep}} = 1$ ) and vary the remaining four parameters systematically on a grid, yielding the best-fit model shown in Fig. 10. The corresponding parameters and uncertainties are listed in column 4 of Tab. 3. In a second step, we test also non-Keplerian velocity fields, yielding the best-fit values listed in column 5 (Fig. 12). The resulting  $\chi^2$ -surfaces are shown in

Fig. 11 and the uncertainties have been derived using the bootstrapping technique.

Our model fits show that the intriguing phase inversion observed at the line center on our longest VLTI baselines (Fig. 10, *3rd row*) can be explained with the phase jumps in the Fourier phase crossing a visibility null. These phase jumps appear in the same spectral channels where we measure visibility minima in the W-shaped visibility profile, indicating that the visibility function of the pure line-emitting geometry transits here from the first to the second visibility lobe (Fig. 10, *2nd and 3rd row*). At the same time, the continuum emission is only marginally resolved, which results in the measured composite line+continuum visibility with a rather high contrast of  $\sim 0.7$ . The phase jumps and visibility minima are a basic property of the Fourier transform of strongly resolved objects and are reproduced very naturally and without finetuning from our kinematical modeling. Therefore, our results do not support the idea outlined by Stefl et al. (2011) that these features might indicate secondary dynamical effects or polar mass outflows.

As best-fit value for the radial intensity index, we yield  $q = -1.6 \pm 0.2$ . Isothermal viscous accretion disk models, such as discussed in Bjorkman & Carciofi (2005), predict a radial disk surface density law  $\Sigma(r) \propto r^{-2}$ , corresponding to  $q = -2$  in the optically thin case. More realistic non-LTE disk models suggest a more shallow surface density profile ( $\Sigma(r) \propto r^{-2 \dots -1}$ ) due to a steep temperature drop in the inner few stellar radii of the disk Carciofi & Bjorkman (2008). Therefore, we conclude that our measured intensity profile is well consistent with these models.

The most significant deviation of our simple kinemat-

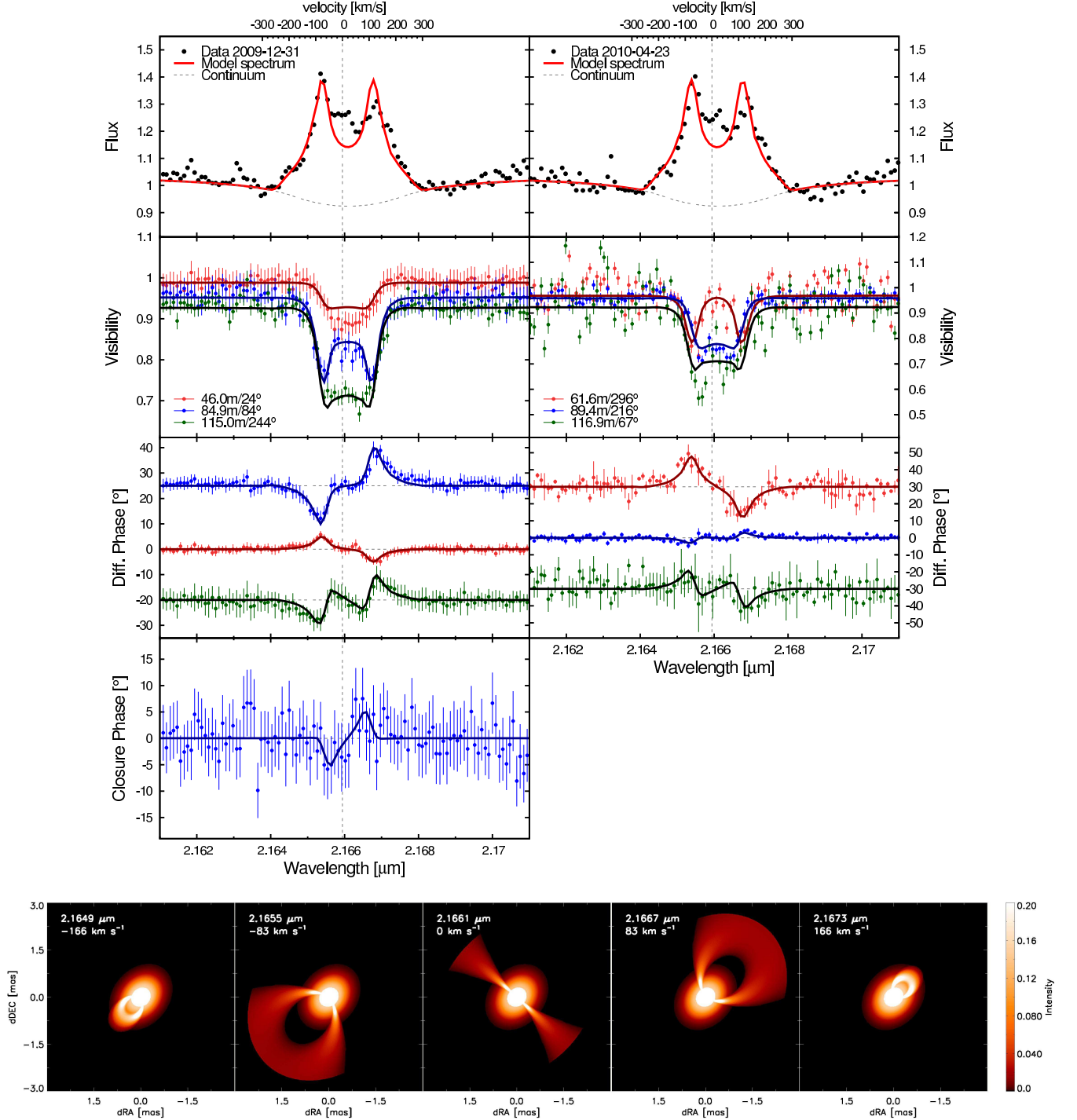


FIG. 10.— *Upper panel:* Comparison of the VLT/AMBER  $\beta$  CMi spectra (*1st row*), visibilities (*2nd row*), DPs (*3rd row*), and CPs (*4th row*) with our Keplerian disk model (solid lines). The three different colors correspond to the three baselines, where the associated projected baseline lengths and PAs are labeled in the 2nd row. *Lower panel:* To illustrate our kinematical model, we show intensity channel maps for five representative wavelengths.



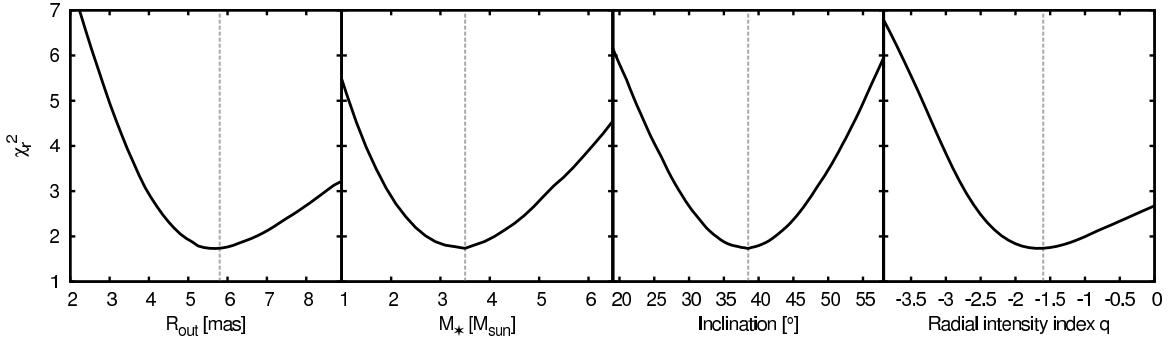


FIG. 11.—  $\chi^2$ -surface derived from our  $\beta$  CMi model grid (Sect. 4.5) around the best-fit solution assuming a Keplerian velocity field (Tab. 3).

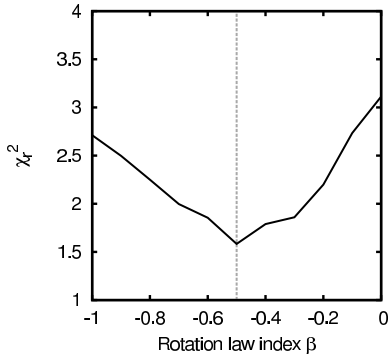


FIG. 12.—  $\chi^2$ -surface for the rotation law index  $\beta$  derived from our  $\beta$  CMi model grid including also non-Keplerian velocity fields (Sect. 4.5).

ical model from the data is in the precise shape of the  $\text{Br}\gamma$ -spectrum. For instance, the measured spectra show a weak emission component in the line center (Fig. 10), which is not reproduced by the model and which might be related either to the presence of a uniformly distributed low-velocity gas halo or, more likely, to radiative transfer effects.

Another deviation between the model and the observation concerns the weak  $V/R$ -asymmetry which can be observed in our VLTI/AMBER  $\text{Br}\gamma$  spectra at both epochs (Fig. 3, *top row*) and which is not reproduced by our axisymmetric kinematical model. Within the 113 days covered by our AMBER observations, no changes in the  $V/R$  asymmetry could be observed, which is consistent with the conclusion by Tycner et al. (2005) and Jones et al. (2011) that the  $\text{H}\alpha$  profile does not show significant long-term variability. It is interesting to compare the average photocenter displacement between the blue- and red-shifted line emission (2.5 mas) with the characteristic size of the  $\text{H}\alpha$ -emitting region (2.13 mas, Tycner et al. 2005), which suggests that  $\text{Br}\gamma$  emerges from a similar spatial region as  $\text{H}\alpha$  and a much more extended region than the  $H$ - and  $K$ -band continuum emission ( $R_{\text{Br}\gamma} \sim R_{\text{H}\alpha} > R_{\text{cont}}$ , Fig. 3), in agreement with the prediction from Carciofi (2010). However, a detailed comparison is difficult due to the non-simultaneity of the different observations.

## 5. DISCUSSION ON $\zeta$ TAURI

$\zeta$  Tau is a particularly well-studied classical Be star which shows a cyclic variability in the flux ratio of the violet- and red-shifted wing of the double-peaked

$\text{H}\alpha$  emission line (Rivinius et al. 2006; Štefl et al. 2009). These  $V/R$  variations exhibit a period of  $\sim 1429$  days and are generally attributed to a global one-armed density oscillation in a Keplerian (or nearly Keplerian) disk (Okazaki 1991).

A kinematic model for the  $\text{Br}\gamma$ -line of some earlier AMBER MR observations on  $\zeta$  Tau was presented by Štefl et al. (2009) and Carciofi et al. (2009) and suggested a Keplerian rotation profile. Our spectro-interferometric observations on  $\zeta$  Tau provide a  $\text{Br}\gamma$  measurement at a new epoch in the  $\text{H}\alpha$   $V/R$ -phase and cover, for the first time, also the wavelength region around the hydrogen Pfund lines. Besides our  $\text{Br}\gamma$  and Pfund series data, we include  $\text{H}\alpha$ -sizes from the literature, namely the GI2T photocenter measurement by Vakili et al. (1998,  $\sim 7 R_*$  or  $\sim 1.33$  mas) and the Mark III and NPOI measurements by Quirrenbach et al. (1994, 1997) and Tycner et al. (2004). It is important to note that the Mark III and NPOI results were based on visibility amplitudes instead of differential phases. Also, a direct comparison with these earlier observations is complicated by the known  $V/R$  variability of  $\zeta$  Tau, although we note that our measurement at phase 0.484 is reasonably close to the  $\text{H}\alpha$ -measurement at phase 0.577, presented by Tycner et al. (2004). We convert the Gaussian FWHM derived by these studies to photocenter displacements by computing the centroid of the corresponding Gaussian brightness distributions. We find that  $\text{Br}\gamma$  emerges

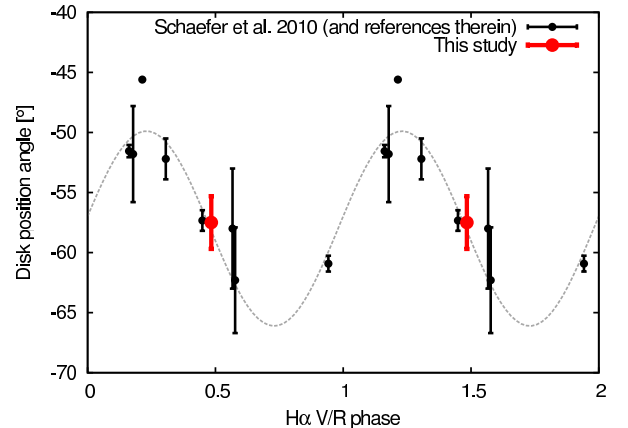


FIG. 13.—  $\zeta$  Tau disk position angle, plotted as function of the  $\text{H}\alpha$   $V/R$  phase, including data presented in Schaefer et al. (2010) and in this study. Our new phase measurement (phase 0.484) is in good agreement with the disk precession law proposed by Schaefer et al.

from a similar spatial region than  $H\alpha$ , but a more extended region than the near-infrared continuum emission, as already found for  $\beta$  CMi (Sect. 4.5). The Pfund emission originates from intermediate stellocentric radii ( $R_{H\alpha} \sim R_{Br\gamma} > R_{Pf} > R_{cont}$ ; Fig. 4).

### 5.1. Signatures of the known one-armed oscillation

Both in  $Br\gamma$  and in the Pf14-Pf22 lines, we detect a double-peaked line profile and clear rotation signatures in the differential phases. All line transitions exhibit a photocenter displacement with a stronger amplitude in the south-eastern (red-shifted) than in the north-western (blue-shifted) lobe (Fig. 4, *middle panel*). Such an asymmetric displacement is consistent with the presence of a one-armed oscillation in the disk (Štefl et al. 2009; Carciofi et al. 2009). CHARA/MIRC observations by Schaefer et al. (2010) showed that the one-armed oscillation pattern can also be observed as an asymmetry in the  $H$ -band continuum emission. Using multi-epoch data, they find that the position angle of the asymmetry is correlated with the spectroscopic variability and precesses around the star with the  $H\alpha$   $V/R$  period. Our observation (2010-01-01) adds an additional epoch for this analysis at the  $V/R$  phase 0.484 (assuming maximum phase at JD=2454505.0 and a period of 1429 days, Schaefer et al. 2010). We find that the spiral density maximum is located towards the south-east of the central star and the measured PA is consistent with the sinusoidal PA modulation proposed by Schaefer et al. (2010), as shown in Fig. 13.

The reported polarization angle for  $\zeta$  Tau ( $32.9 - 33.1^\circ$ , Quirrenbach et al. 1997;  $35 \pm 4^\circ$ , Ghosh et al. 1999;  $32 - 33.5^\circ$ , McDavid 1999) is in excellent agreement with the rotation axis position angle ( $35.8 \pm 2.0$ ) found by our spectro-interferometric observations.

### 5.2. Probing the disk excitation structure using multi-transition spectro-interferometry

In order to test whether the measured differences in the stellocentric emitting radius of the  $Br\gamma$  and Pfund transitions are consistent with the expected excitation structure in the disk, we construct a simple radiative transfer model assuming local thermodynamic equilibrium (LTE). We assume that the line-emitting mate-

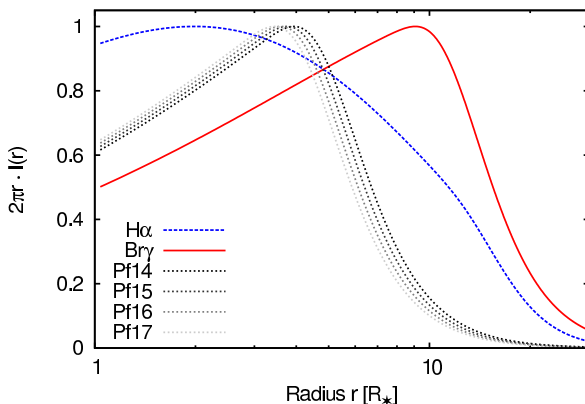


FIG. 14.— Radial intensity profile, as computed with our LTE model for different line transitions (Sect. 5.2). The intensity has been weighted by the emitting area and normalized to the peak intensity.

rial is located in an equatorial disk which extends outwards from the stellar radius with a constant vertical density per unit volume and a half-opening angle of  $\Theta = 5^\circ$ . Then, we integrate for each radius  $r$  the optical depth  $\tau_\nu$  in vertical direction, assuming hydrogen under LTE conditions and an isothermal temperature distributions. The number density for the different excitation levels and ionization stages is computed using the Saha and Boltzmann equation assuming a Gaussian line profile with thermal line broadening (Wilson et al. 2009). The Einstein coefficients for the different hydrogen transitions are estimated using the series expansion published by Omidvar & McAllister (1995). The radial temperature and surface density profile are parameterized with  $T(r) = T_{\text{eff}}(r/R_\star)^{-0.5}$  (Stee & de Araujo 1994) and  $\Sigma(r) = \Sigma_0(r/R_\star)^{-2}$  with  $\Sigma_0 = 2.1 \text{ g cm}^{-2}$  (Carciofi et al. 2009). For the stellar temperature, equatorial stellar radius, and distance, we assume  $T_{\text{eff}} = 19370 \text{ K}$ ,  $R_\star = 7.7 R_\odot$ , and  $d = 126 \text{ pc}$ , respectively (Carciofi et al. 2009).

Using the radiative transfer equation for LTE conditions ( $I_\nu(r) = B_\nu(T)(1 - e^{-\tau_\nu(r)})$ ), we compute the emitted intensity per unit area as function of radius  $r$  in the disk, where  $B_\nu$  is the Planck spectrum for temperature  $T$  at frequency  $\nu$ . From the radial intensity profiles (Fig. 14) we compute the centroid of the brightness distribution, which is then compared to the stellocentric emission radius measured in the different line transitions. Given that our model assumes a simplified vertical density structure and does not include inclination effects, we do not aim to match the absolute sizes of the emitting region in all line transitions, but focus instead on the *relative* sizes. For this step, we normalize both the measured and the model photocenter offsets relative to  $Br\gamma$ . The comparison between the  $H\alpha$ ,  $Br\gamma$ , and Pf14-17 relative sizes and our LTE model is shown in Fig. 15. We find that we can reproduce important observational features, in particular that

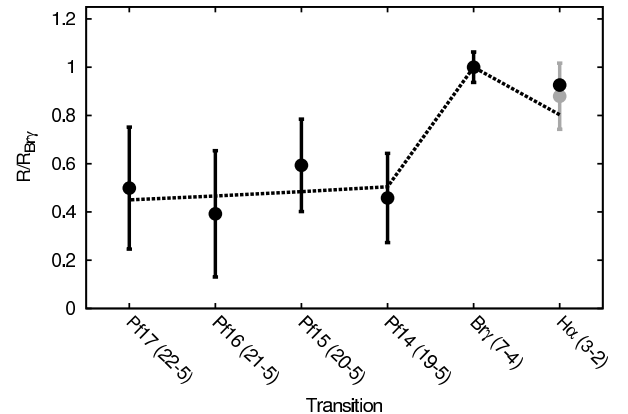


FIG. 15.— Origin of the  $Br\gamma$  and Pfund line emission from  $\zeta$  Tau, as measured by the length of the photocenter displacement vector, averaged over all velocity channels for a given line transition (the error bars represent the standard deviation in the different velocity channels). In addition, we include the  $H\alpha$ -line photocenter displacement by Vakili et al. (1998, black data point) and the  $H\alpha$  size estimates from Quirrenbach et al. (1994), Quirrenbach et al. (1997), Tycner et al. (2004), which we converted from Gaussian FWHM to the corresponding photocenters (grey data point). The dashed line shows the prediction from our LTE excitation model (Sect. 5.2).

- a) Br $\gamma$  originates from a similar spatial region in the disk as H $\alpha$  ( $R_{\text{Br}\gamma} \sim R_{\text{H}\alpha}$ ). This result is also in agreement with the predictions from more sophisticated non-LTE radiative transfer computations, such as made by Carciofi (2010). In order to better characterize the differences between the Br $\gamma$  and H $\alpha$ -emitting region, contemporaneous observations in these two wavelength bands will be required.
- b) Br $\gamma$  originates from a more extended region than the Pfund lines ( $R_{\text{Br}\gamma} > R_{\text{Pf}}$ ). Computing the weighted average of the measurements in the individual transitions, we yield  $R_{\text{Pf}}/R_{\text{Br}\gamma} = 0.48 \pm 0.12$ . For the Pfund lines, we are not aware of earlier predictions obtained with non-LTE radiative transfer codes.

Our observational results confirm the finding from Pott et al. (2010), which found  $R_{\text{Br}\gamma} > R_{\text{Pf}}$  on the classical Be star 48 Lib, and consolidates their suggestion that the measured size differences can already be explained with the expected optical depth differences between these line transitions. We encourage theoreticians and modelers to employ their more sophisticated non-LTE radiative transfer codes in order to test the influence of inclination and non-LTE effects, although, based on the results from Iwamatsu & Hirata (2008, e.g. Fig. 3), we expect no significant departure from LTE for the inner disk regions and the high transitions traced by our observations. Future multi-transition spectro-interferometric observations with improved  $uv$ -coverage might also measure the precise radial intensity profile in a model-independent fashion. Together with sophisticated radiative transfer simulations, these observations will reveal the excitation structure of the disk and constrain parameters such as the temperature profile and the vertical disk structure, which are currently difficult to access.

## 6. CONCLUSIONS

Using CHARA and VLTI interferometry, our study combined high angular resolution (with baseline lengths up to 330 m) with kinematical information obtained at high spectral dispersion, yielding direct observational constraints on the gas distribution, excitation structure, and kinematics of the disks around two classical Be stars. Using a model-independent photocenter analysis method we derived the disk rotation axis for the prototypical objects  $\beta$  CMi and  $\zeta$  Tau and spatially and spectrally resolved the disk rotation profile on scales of a few stellar radii. For both objects, we find that the determined gas rotation plane agrees well with the orientation of the continuum-emitting disk as resolved by CHARA, although there is also clear evidence for substructure in the disk around  $\zeta$  Tau, revealing a one-armed oscillation, as indicated by different displacement amplitudes in the blue- and red-shifted line wing.

Using our data set on  $\beta$  CMi we constructed a position-velocity diagram, which can be interpreted using the well-established procedures from radio interferometry, but probes the milli-arcsecond scale position displacements resulting from the rotating disk around this Be star. From our kinematical constraints, we derive the dynamical mass of the central star to  $3.5 \pm 0.2$ , which

is in excellent agreement with earlier spectroscopic studies (Saio et al. 2007). The inclination of the system is  $38.5 \pm 1^\circ$ , as determined with our CHARA continuum and VLTI line observations. As shown with our detailed kinematical modeling, the rotation law is Keplerian ( $\beta = -0.5 \pm 0.1$ ) and we do not have to include an expanding velocity component in order to explain our data. Furthermore, our kinematical model allowed us to identify the origin of the phase inversion, which has now been observed in the differential phases in five out of eight Be stars. These phase jumps correspond to the transition from the first to the second visibility lobe, removing the necessity for speculations beyond the canonical star+disk paradigm (Steff et al. 2011).

For  $\zeta$  Tau, we obtained spectro-interferometric observations covering simultaneously the Br $\gamma$  and at least nine transitions from the Pfund line series. For all transitions, we detect a significantly stronger photocenter displacement in the red-shifted line wing than in the blue-shifted line wing, tracing the one-armed oscillation which has been deduced for the  $\zeta$  Tau disk before. Comparing the photocenter displacement in the different line transitions, we find that the Pfund, Brackett, and Balmer lines originate from different stellocentric emitting regions ( $R_{\text{cont}} < R_{\text{Pf}} < R_{\text{Br}\gamma} \sim R_{\text{H}\alpha}$ ), which we can reproduce qualitatively with a simple LTE line radiative transfer model. More work, including non-LTE radiative transfer modeling, will be required in order to derive quantitative constraints.

By detecting a purely Keplerian velocity field, our observations are inconsistent with disk-formation mechanisms incorporating a strong outflowing velocity component, such as the wind compression scenario, which predicts a strong radial velocity component comparable to the escape velocity (Bjorkman & Cassinelli 1993). On the other hand, our kinematical constraints, as well as the measured hydrogen line intensity profiles (with a radial power law index  $q = -1.6 \pm 0.2$ ) are consistent with the predictions from Keplerian viscous accretion disk models (Lee et al. 1991). As shown by Kato (1983) and discussed in various reviews (e.g. Carciofi 2010), viscous Keplerian disks are also able to produce one-armed density oscillations, such as detected for  $\zeta$  Tau.

Considering that our kinematic constraints have been obtained using a very limited number of individual measurements (2 pointings on  $\beta$  CMi, 1 pointing on  $\zeta$  Tau), our study also illustrates the high effectiveness achievable with spectro-interferometry, in particular if a very high spectral resolution is employed or several line transitions are observed.

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